

Local phase-space structure of dark-matter and direct detection experiments

Based on
Vogelsberger, Helmi & Aquarius coll.
(MNRAS, 2009)



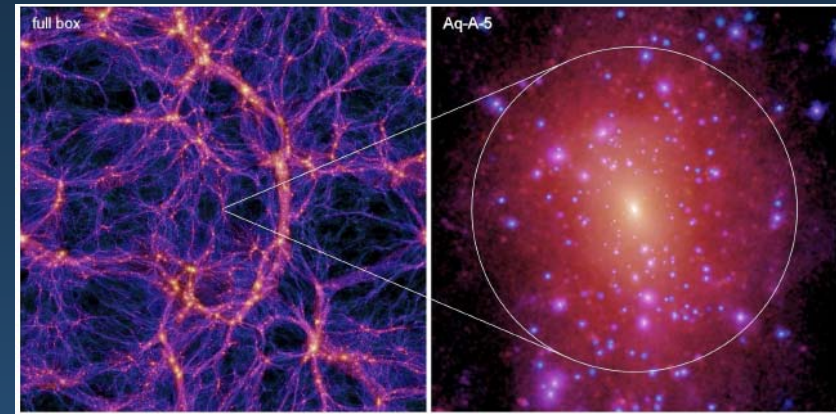
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institute

Direct detection & Aquarius

- Set of very high-resolution simulations of a MW-like dark matter halo

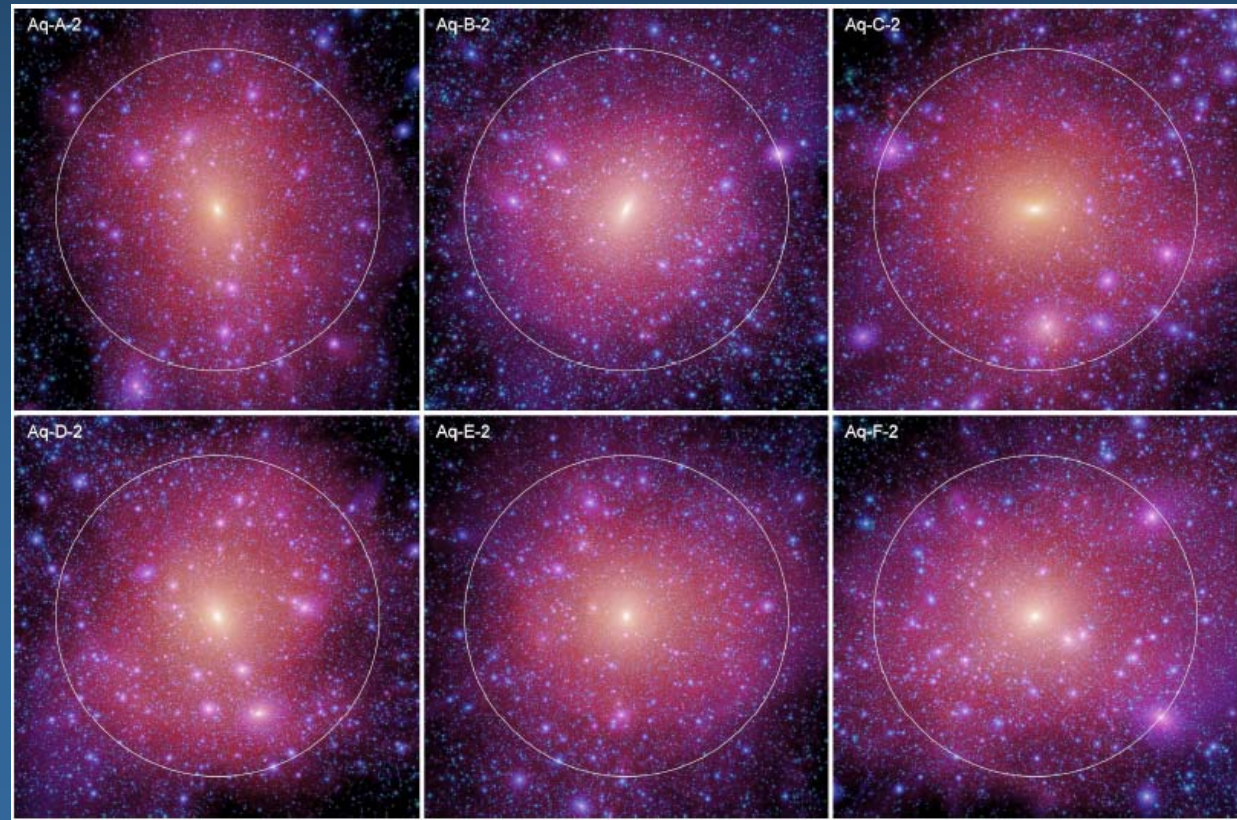


Aquarius A:

- 5 different resolutions
- Ideal for convergence tests

Aquarius B-F:

- 6 different dark halos
- Individual formation histories
- Test scatter in properties



Velocity distribution near the "Sun"

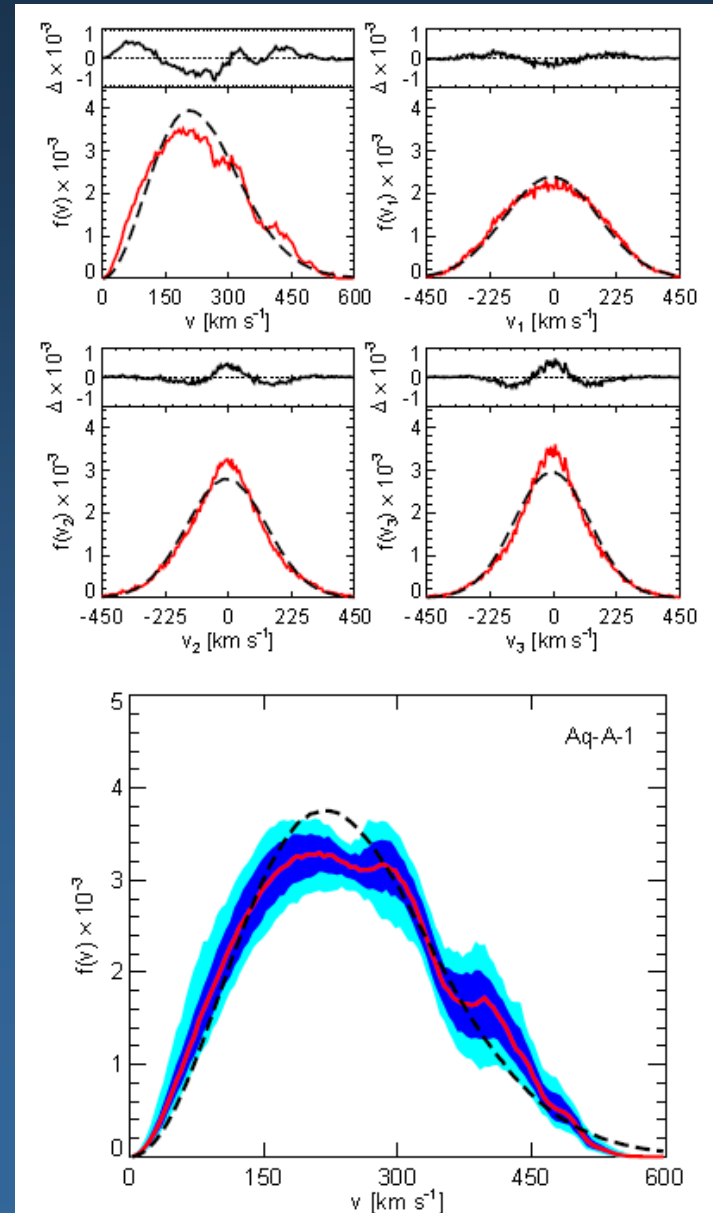
- Velocity distribution of particles in small volumes on the "solar circle"

- No v-component follows a Gaussian

- Distribution of the modulus of velocity

- Does not correspond to Gaussian
- shows bumps

- Bumps are visible in all volumes (at the same velocity value) and in all experiments



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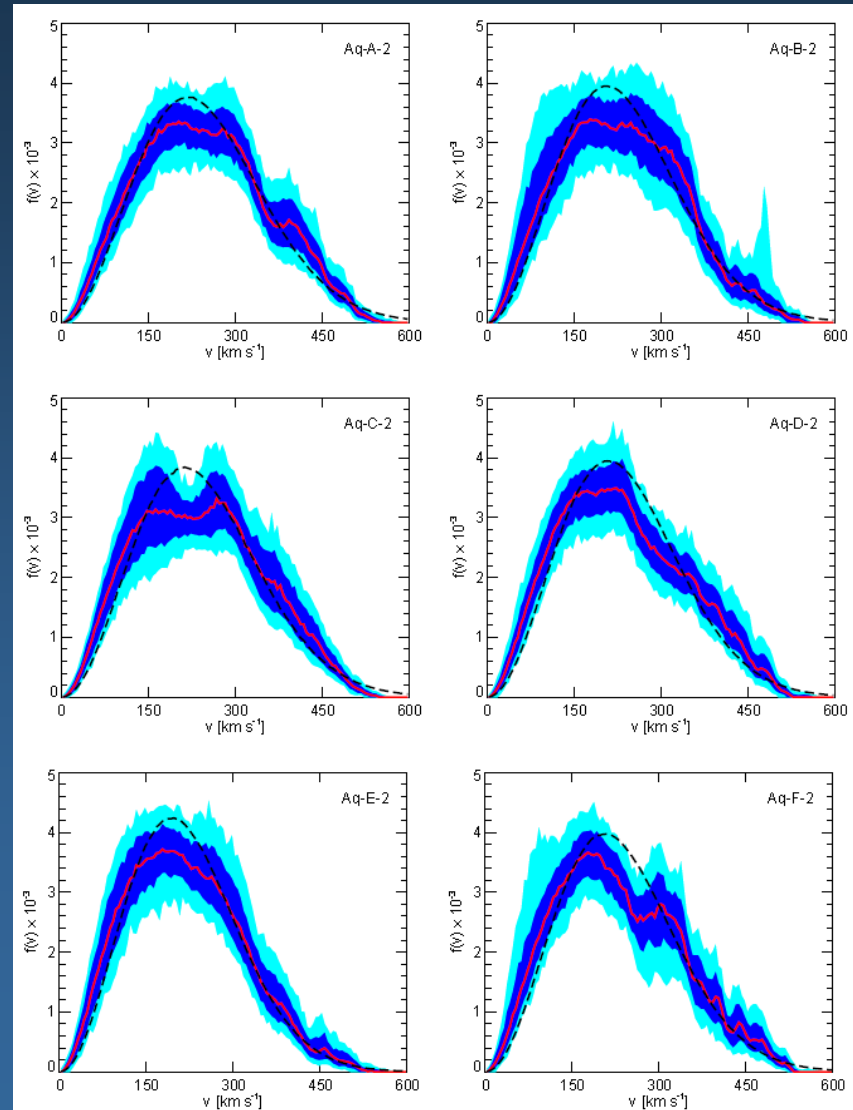
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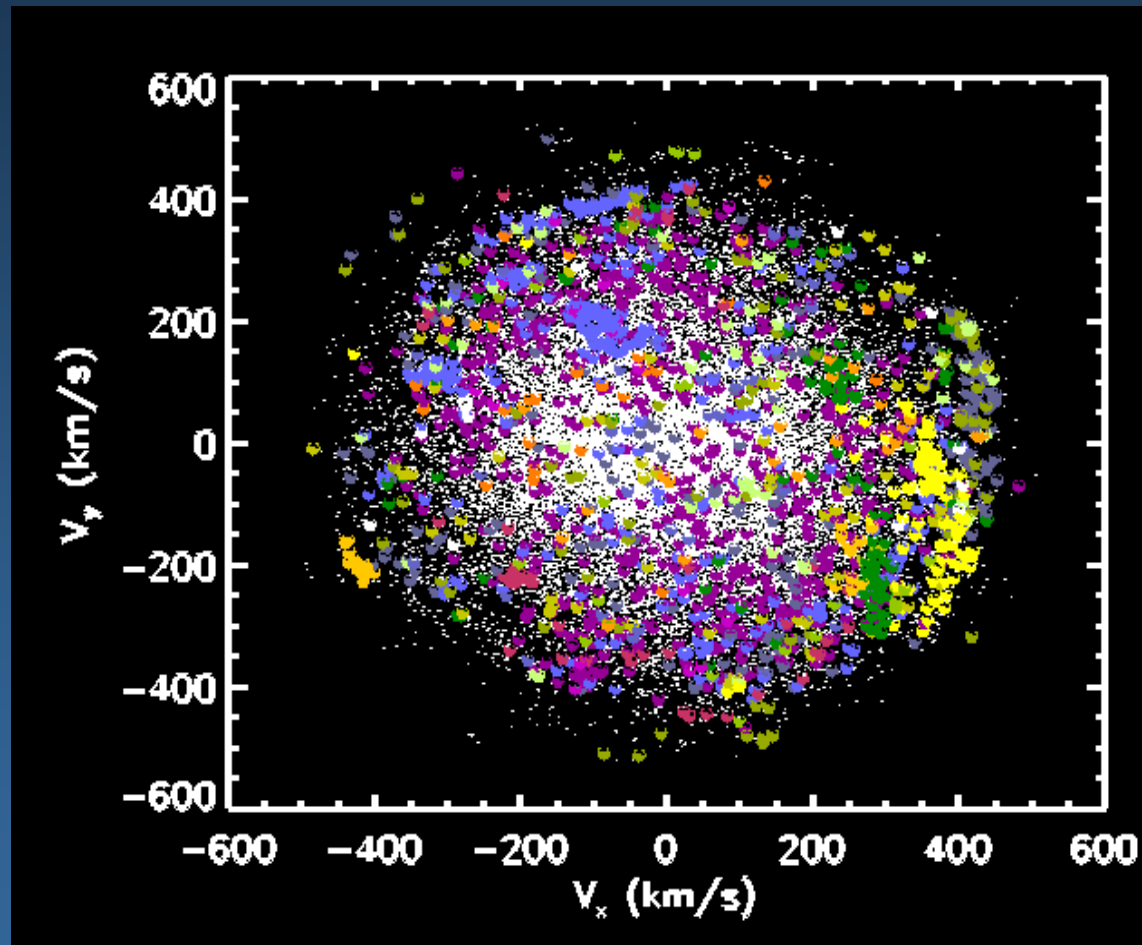
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Aquarius: dark matter near the Sun

- Streams from past mergers clearly resolved the first time near Sun
- No massive (dominant) streams near the "Sun"
- 4×10^4 particles in volume
 - 27 halos contribute at least 10 particles (0.025% of the total)
 - Most prominent streams have ~ 100 particles (0.25% of the total)



Vogelsberger, AH et al. 2009

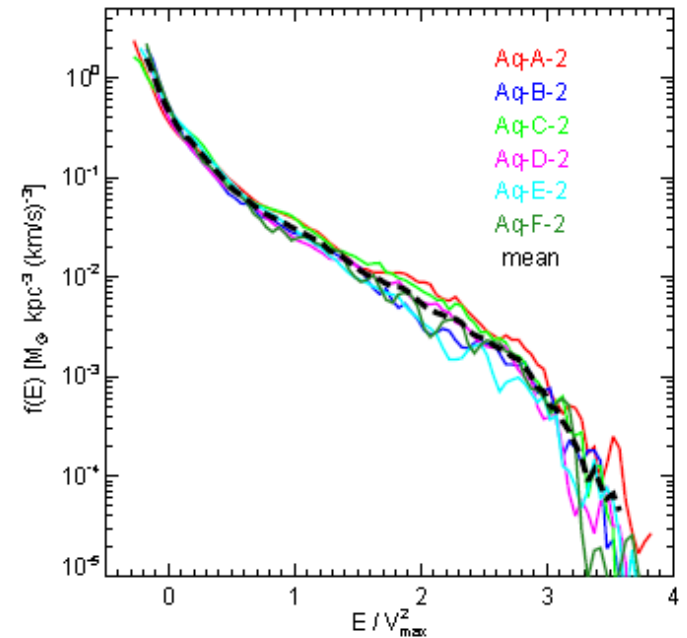
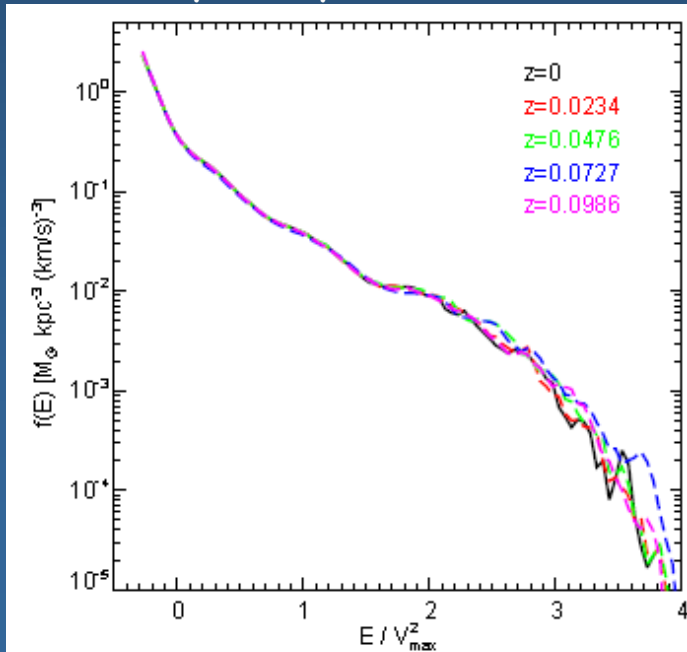
Phase-space distribution $f(E)$

- Bumps in $f(v)$ must be present also in energy space:

$$E = \frac{1}{2} v^2 + \Phi(r)$$

- $f(E)$ shows wiggles

- in all resolutions for Aq-A (in the same location \rightarrow convergence)
- all experiments (must be related to different histories)
- Even for the highest binding energies (typically associated to material accreted very early on)



Phase-space distribution $f(E)$

- Wiggles can be linked to mergers
 - As early as $z = 7$ (> 12.5 Gyr ago)
- less wiggles \rightarrow quiescent history
- many wiggles \rightarrow recent major merger event
- If we could measure $f(E) \rightarrow$ reconstruct history of dark halo!

